H and He burning studied at Gran Sasso

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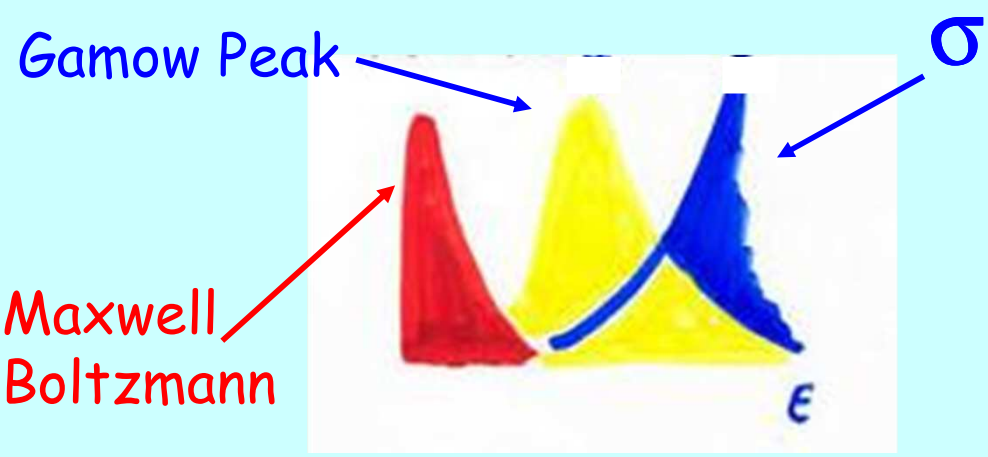
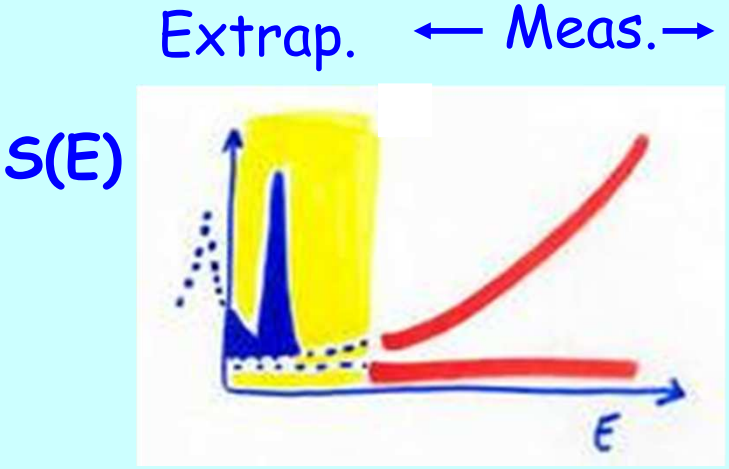
- Stellar Energy+Nucleosynthesis
- Hydrogen + He Burning
- $\sigma(E_{star})$ with $E_{star} << E_{Coulomb}$

$$\sigma(E) = S(E)\, e^{-2\pi\eta}\, E^{-1}$$

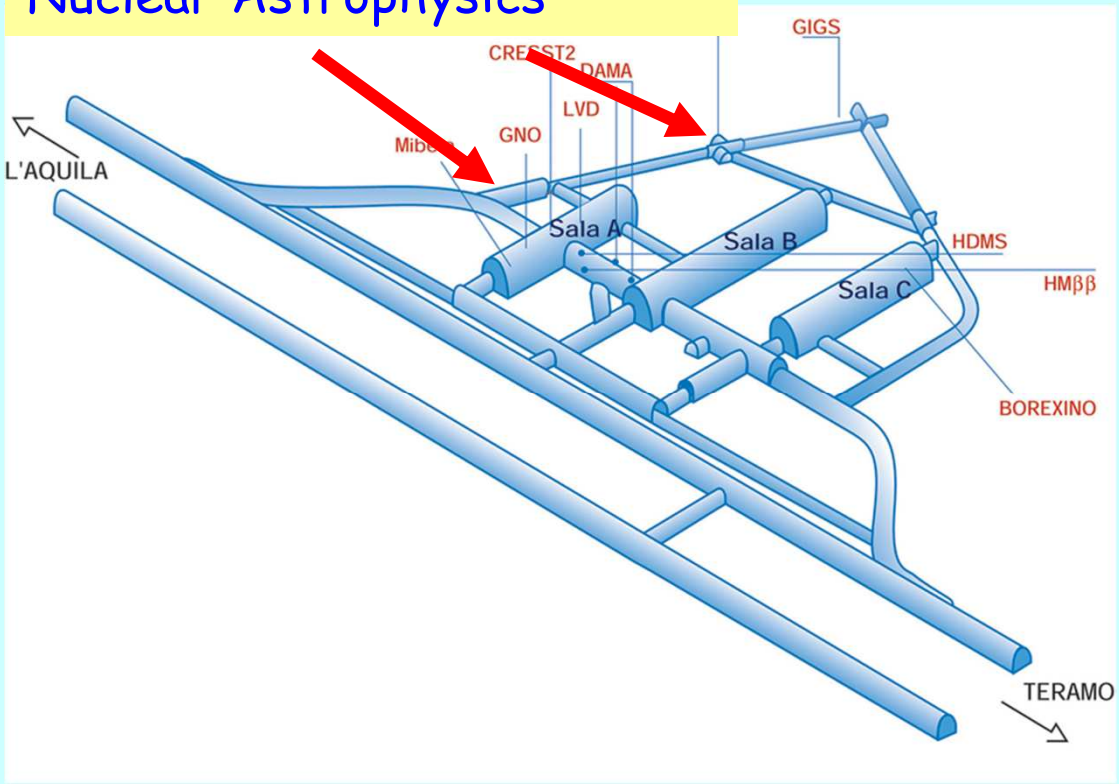
$$2\pi\eta = 31.29\, Z_1 Z_2 \sqrt{\mu/E} \qquad \mu = m_1 m_2 / (m_1+m_2)$$

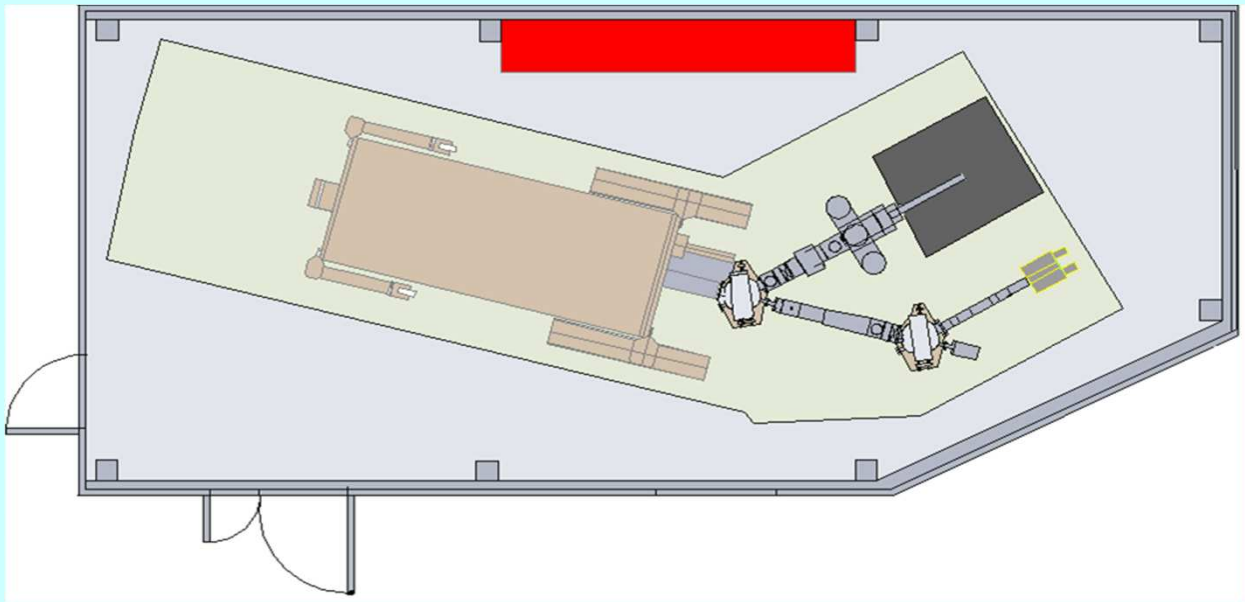
$$\text{Reaction Rate(star)} \div \int \Phi(E)\, \sigma(E)\, dE$$





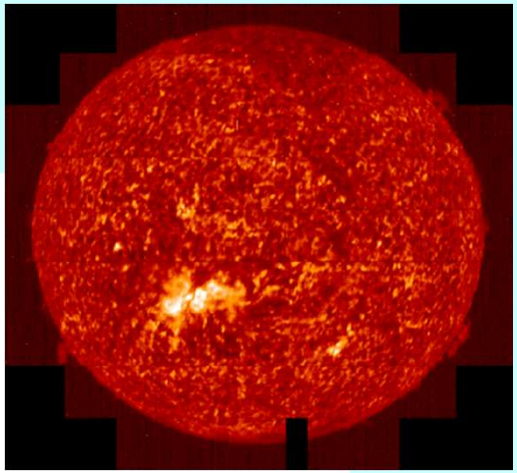
Laboratory for Underground Nuclear Astrophysics





Beam: H,He
Voltage Range :50-400 kV
Output Current: ~1 mA
Absolute Energy error
±300 eV
Beam energy spread:
<100 eV
Long term stability (1 h) :
5 eV
Terminal Voltage ripple:
5 Vpp Ge detector

Hydrogen burning in the Sun @15*10^6 degrees:

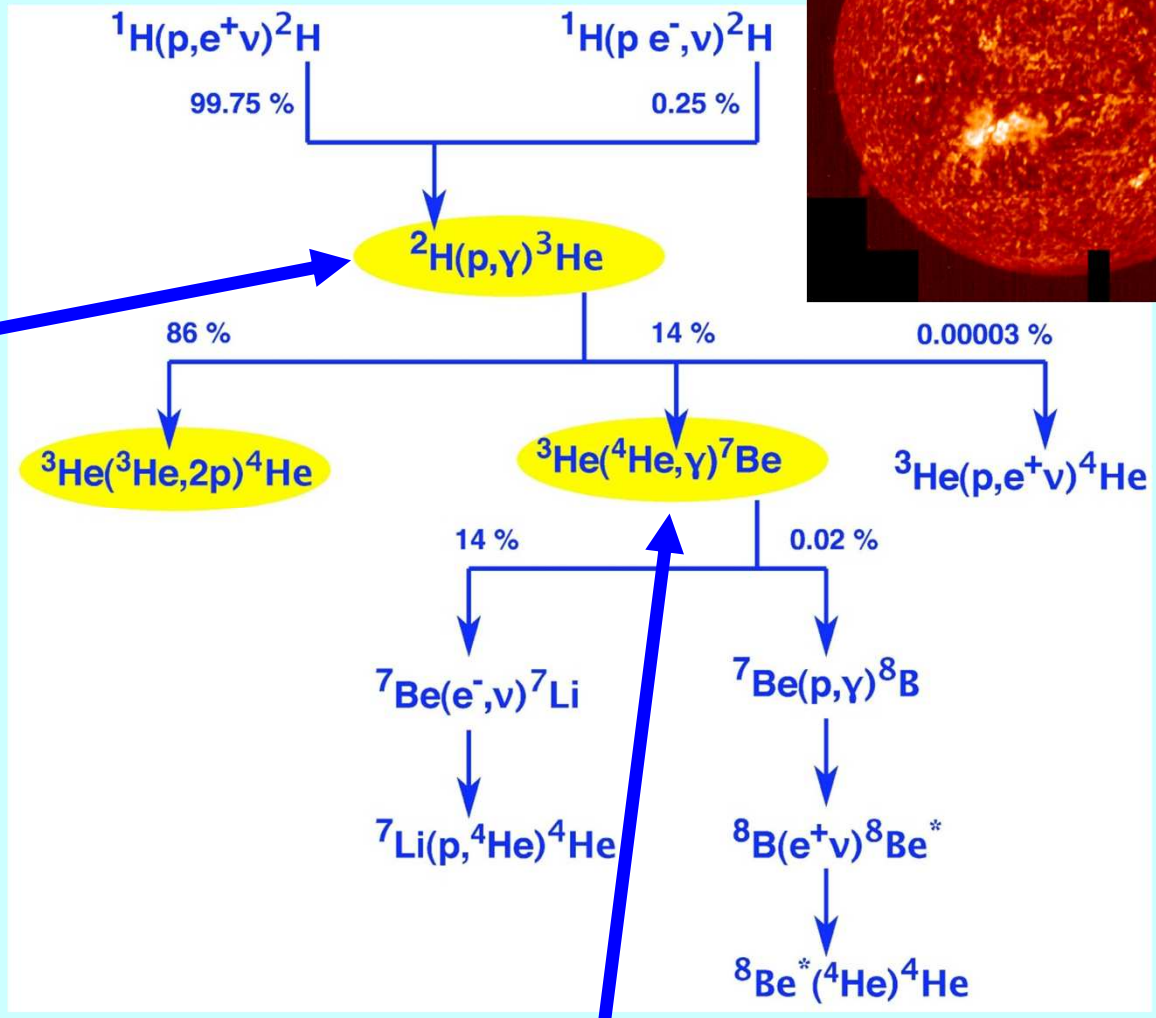


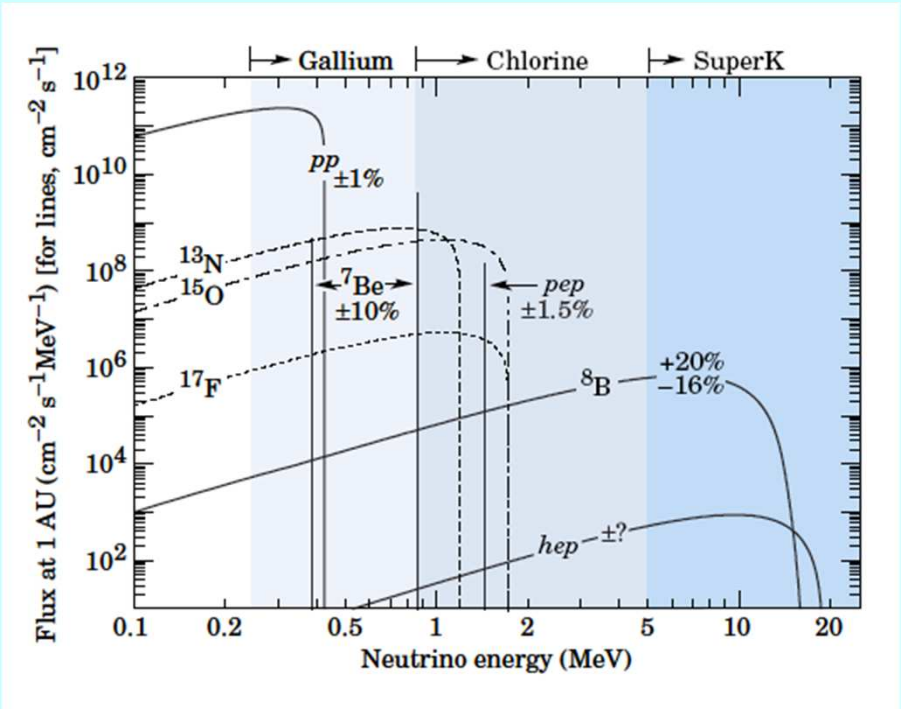
6*10^{11} kg/s of H ⟶ He

+0.7% M_H ⟶ E

^{2}H burning in proto-stars @10^6 degrees

Resonance?

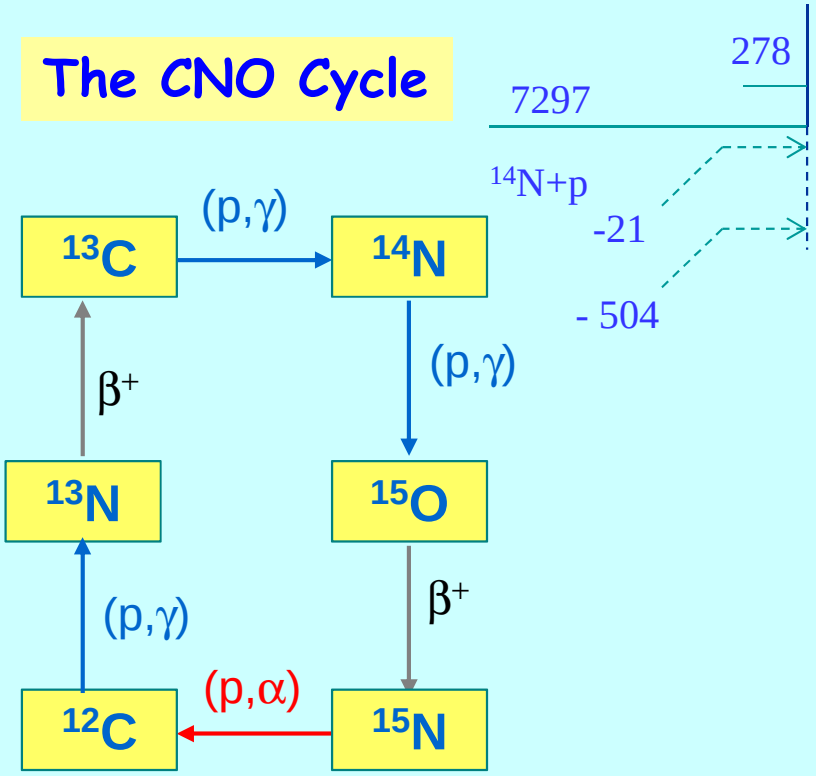


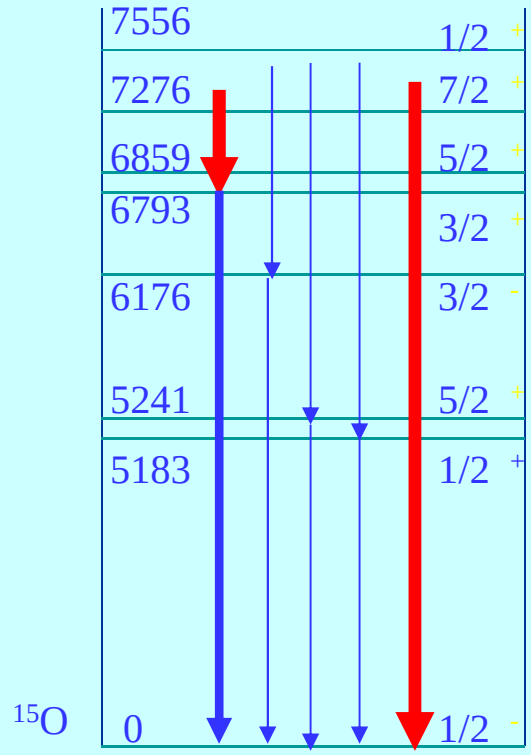


activation=prompt gamma

σ at low energy with 4% error

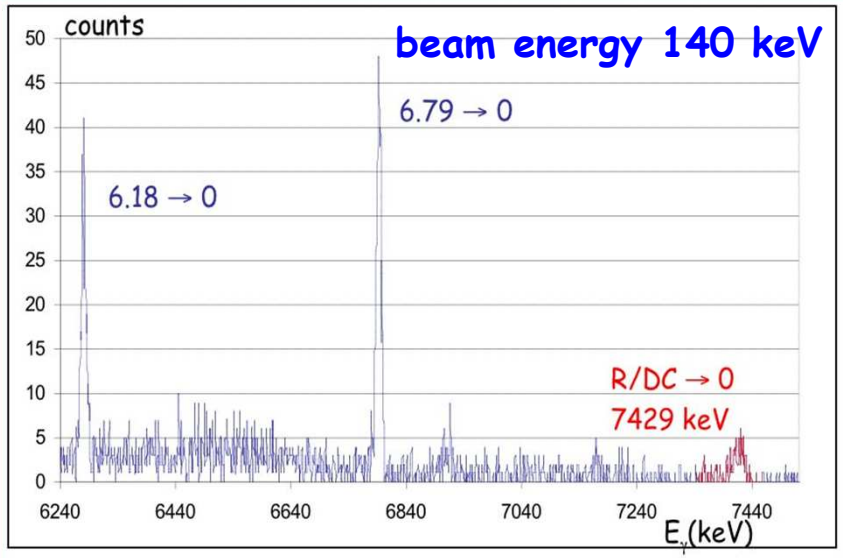
The CNO Cycle



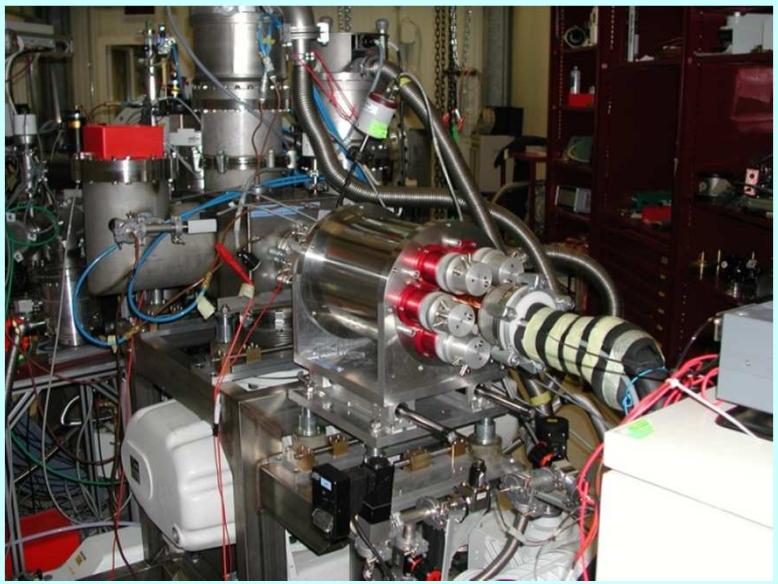


^{14}N(p,γ)^{15}O

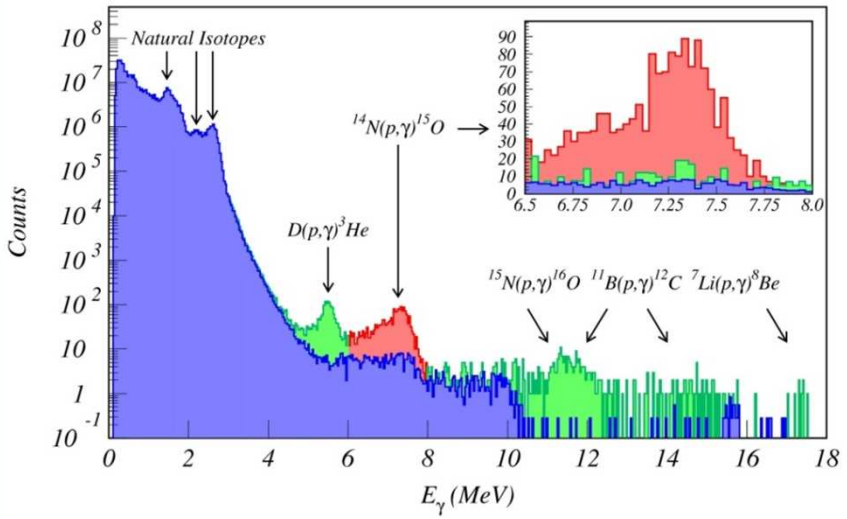
"High" energy: solid target + HpGe

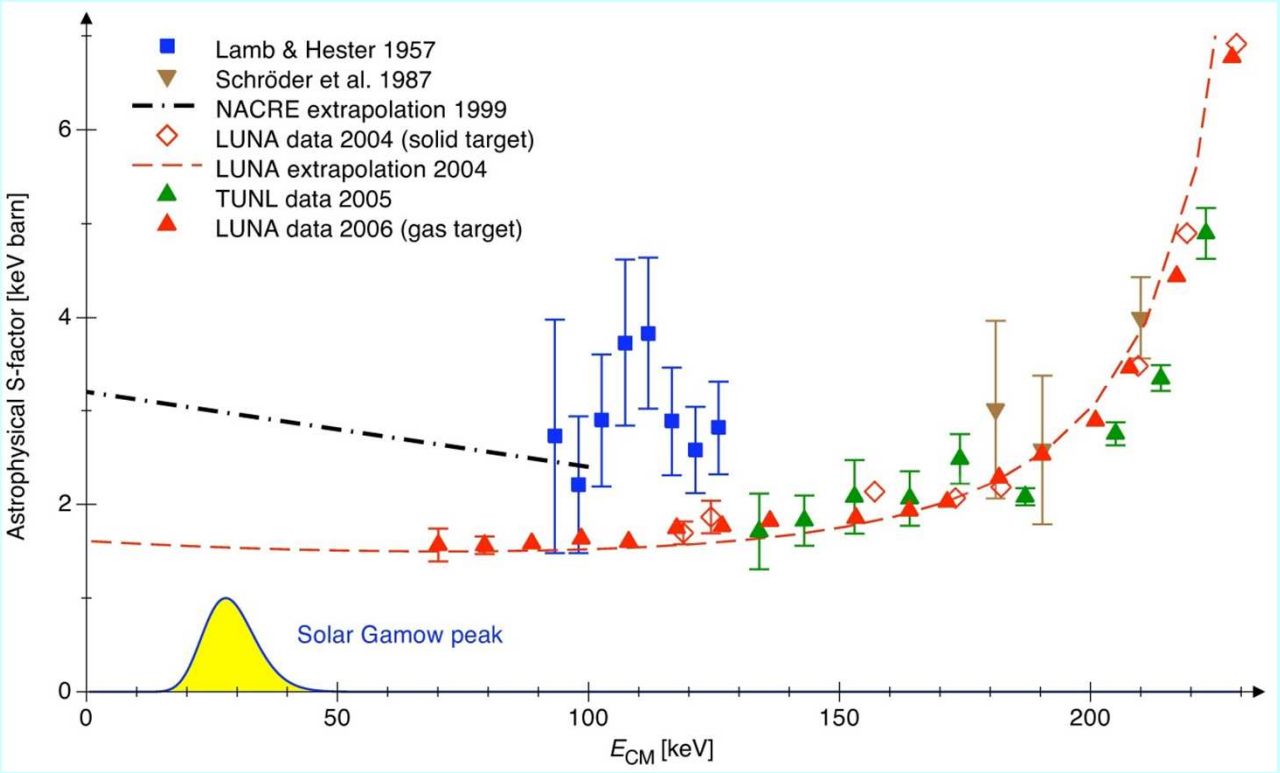


Low energy: gas target + BGO



beam energy 90 keV





$S_1(0)=1.57\pm0.13$ keV b

* $\frac{1}{2}\nu_{cno}$ from the Sun
* Globular Cluster age +1Gy
* more C at the surface of AGB

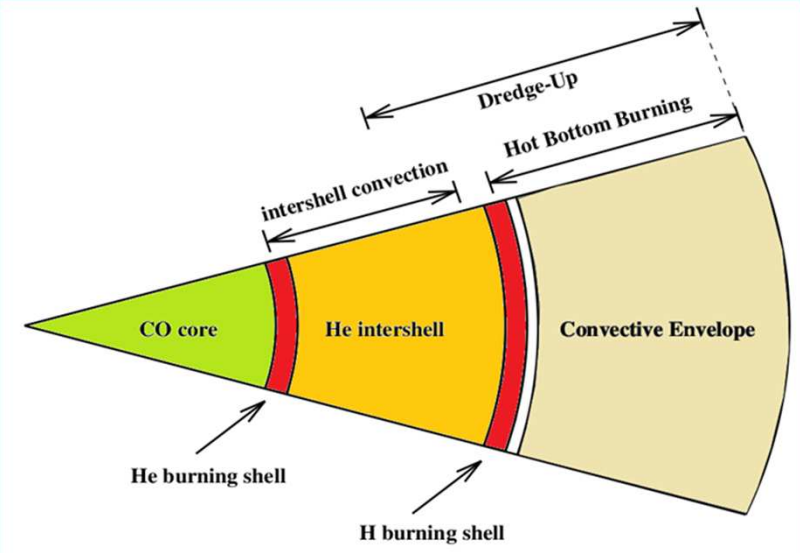
Solar composition problem:
Z/X ~0.024 ⟶ ~ 0.018
SSM predictions disagree with Helioseismology results

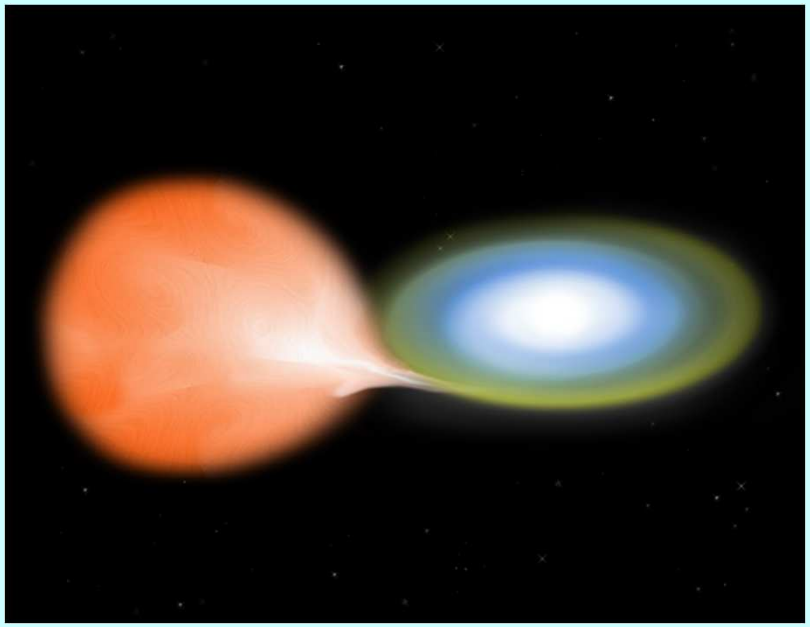
$\nu_{cno}=f(Z,S_{14})$, ~30% decrease from high to low metallicity

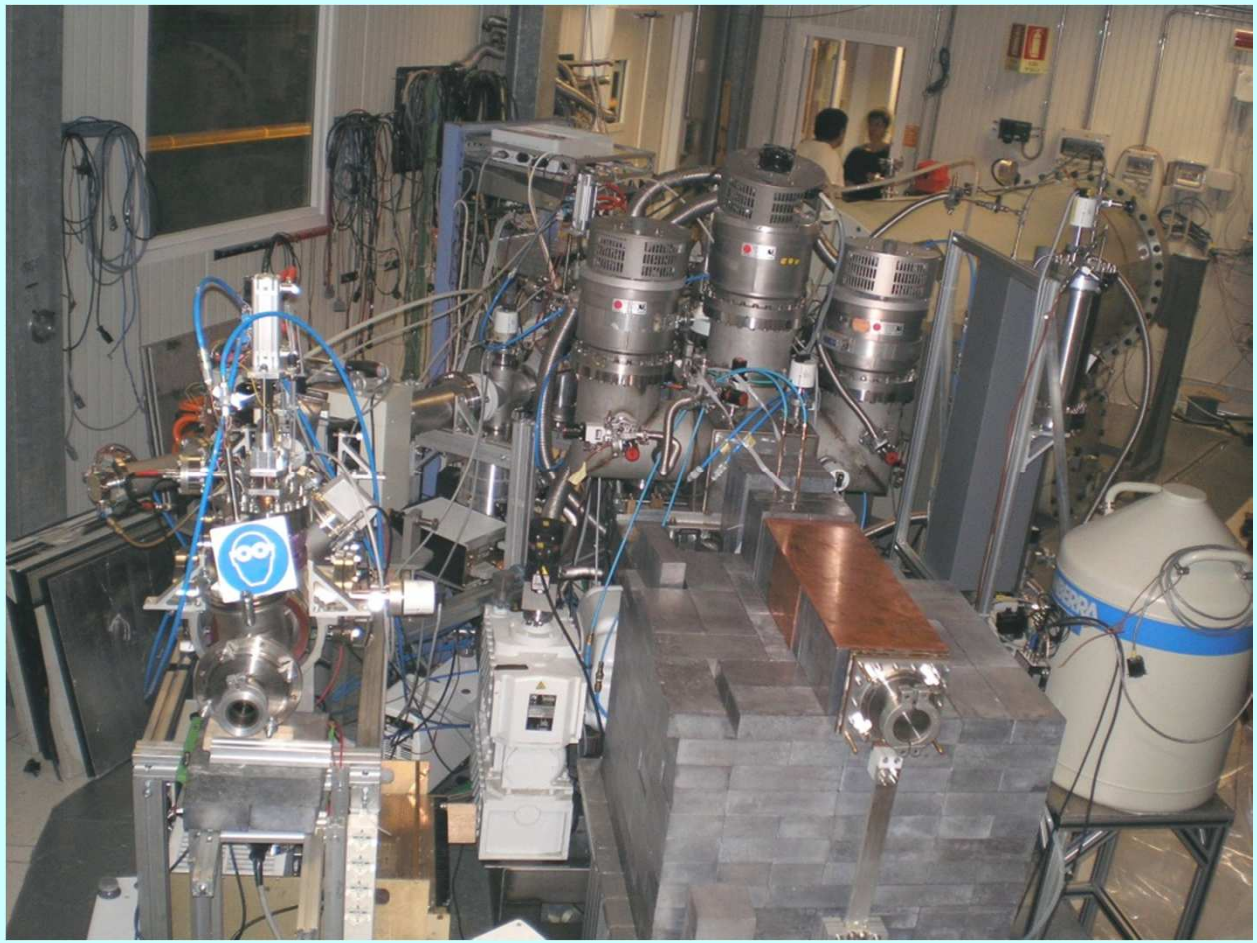
From a measurement of ν_{cno} from the Sun

⟶ Metallicity of the Sun core (C+N)
Photosphere and core metallicity equal?

LUNA beyond the Sun: isotope production in the hydrogen burning shell of AGB stars (~30-100 T_6), Nova nucleosynthesis (~100-400 T_6) and BBN







LUNA beyond the Sun: isotope production in the hydrogen burning shell of AGB stars (~30-100 T_6), Nova nucleosynthesis (~100-400 T_6) and BBN

$^{15}N(p,\gamma)^{16}O$ Q=12.13 MeV

$^{25}Mg(p,\gamma)^{26}Al$ Q=6.3 MeV

$^{17}O(p,\gamma)^{18}F$ Q=5.6 MeV

$^{2}H(\alpha,\gamma)^{6}Li$ Q=1.47 MeV

$^{17}O(p,\alpha)^{14}N$ Q=1.2 MeV

$^{22}Ne(p,\gamma)^{23}Na$ Q=8.8 MeV

$^{18}O(p,\gamma)^{19}F$ Q=8.0 MeV

$^{18}O(p,\alpha)^{15}N$ Q=4.0 MeV

$^{23}Na(p,\gamma)^{24}Mg$ Q=11.7 MeV

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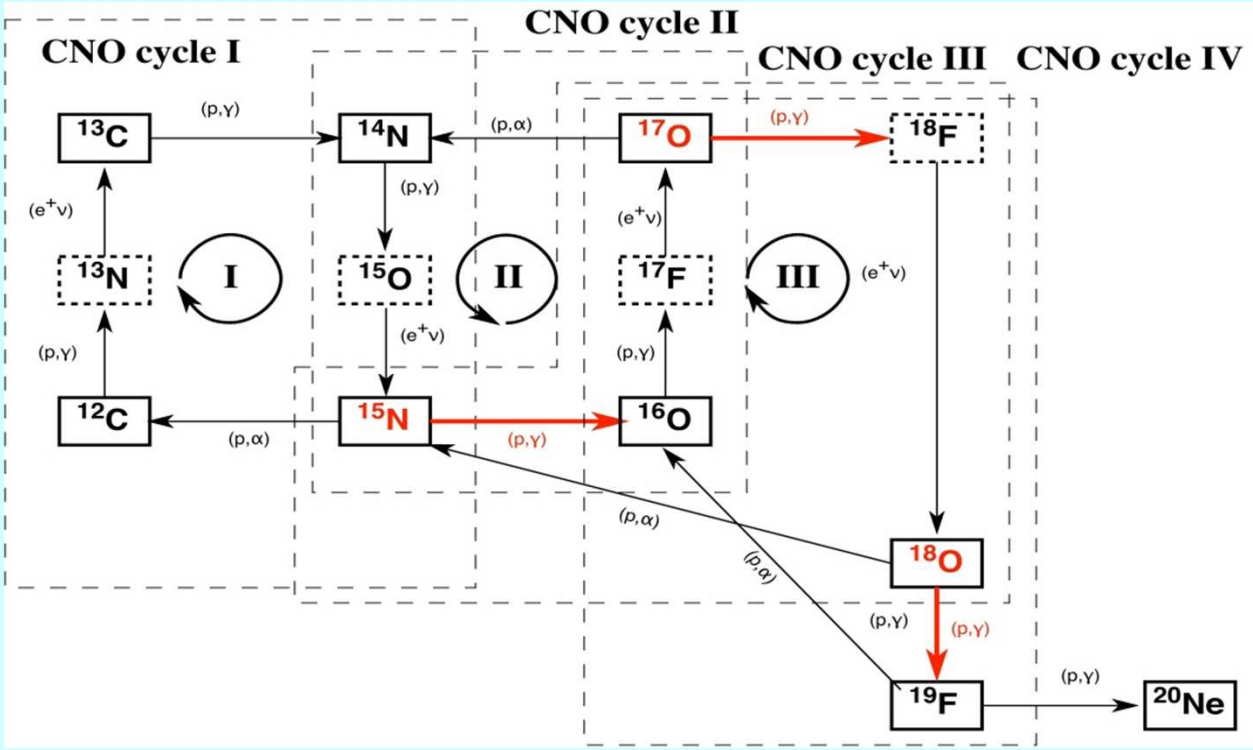

LUNA beyond the Hydrogen burning: 3.5 MV accelerator mainly devoted to **Helium-Burning** (in stars: ~100 T_6, ~10^5 gr/cm^3)

$^{12}C(\alpha,\gamma)^{16}O$ the most important reaction of nuclear astrophysics: production of the elements heavier than A=16, star evolution from He burning to the explosive phase (core collapse and thermonuclear SN) and ratio C/O

Sources of the neutrons responsible for the S-process: 50% of the elements beyond Iron

$^{13}C(\alpha,n)^{16}O$: isotopes with A≥90 during AGB phase of low mass stars

$^{22}Ne(\alpha,n)^{25}Mg$: isotopes with A<90 during He and C burning in massive stars

(α,γ) on ^{3}He, ^{14}N, ^{15}N, ^{18}O……

LUNA-MV partially financed by MIUR

'Starting the LUNA-MV Collaboration', LNGS, February 6th-8th 2013
First LUNA-MV meeting, September 30th 2013

- $^{3}He\ (^{3}He,2p)^{4}He$: σ down to 16 keV

  no resonance within the solar Gamow Peak

- $^{3}He(\alpha,\gamma)^{7}Be$: $^{7}Be \approx$ prompt γ

  Cross section measured with 4% error

- $^{14}N(p,\gamma)^{15}O$: σ down to 70 keV

  ν_{cno} reduced by ~ 2 with 8% error → Sun core metallicity

  Globular cluster age increased by 0.7-1 Gy

  More carbon at the surface of AGB stars

- $^{15}N(p,\gamma)^{16}O$: σ down to 70 keV, reduced by ~ 2

- $^{25}Mg(p,\gamma)^{26}Al$: first measurement of the 92 keV resonance,

  strength $\omega\gamma=(2.9\pm0.6)\times 10^{-10}$ eV

- $^{17}O(p,\gamma)^{18}F$: rate uncertainty @ Novae temperature reduced to 5%

  → uncertainty on ^{18}O, ^{18}F and ^{19}F less than 10% (from 40-50%)

- Future: Hydrogen and **Helium burning (3.5 MV accelerator)**

LUNA Collaboration

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